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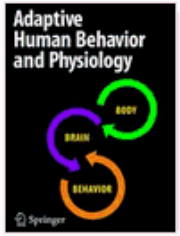

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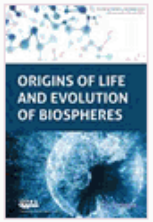

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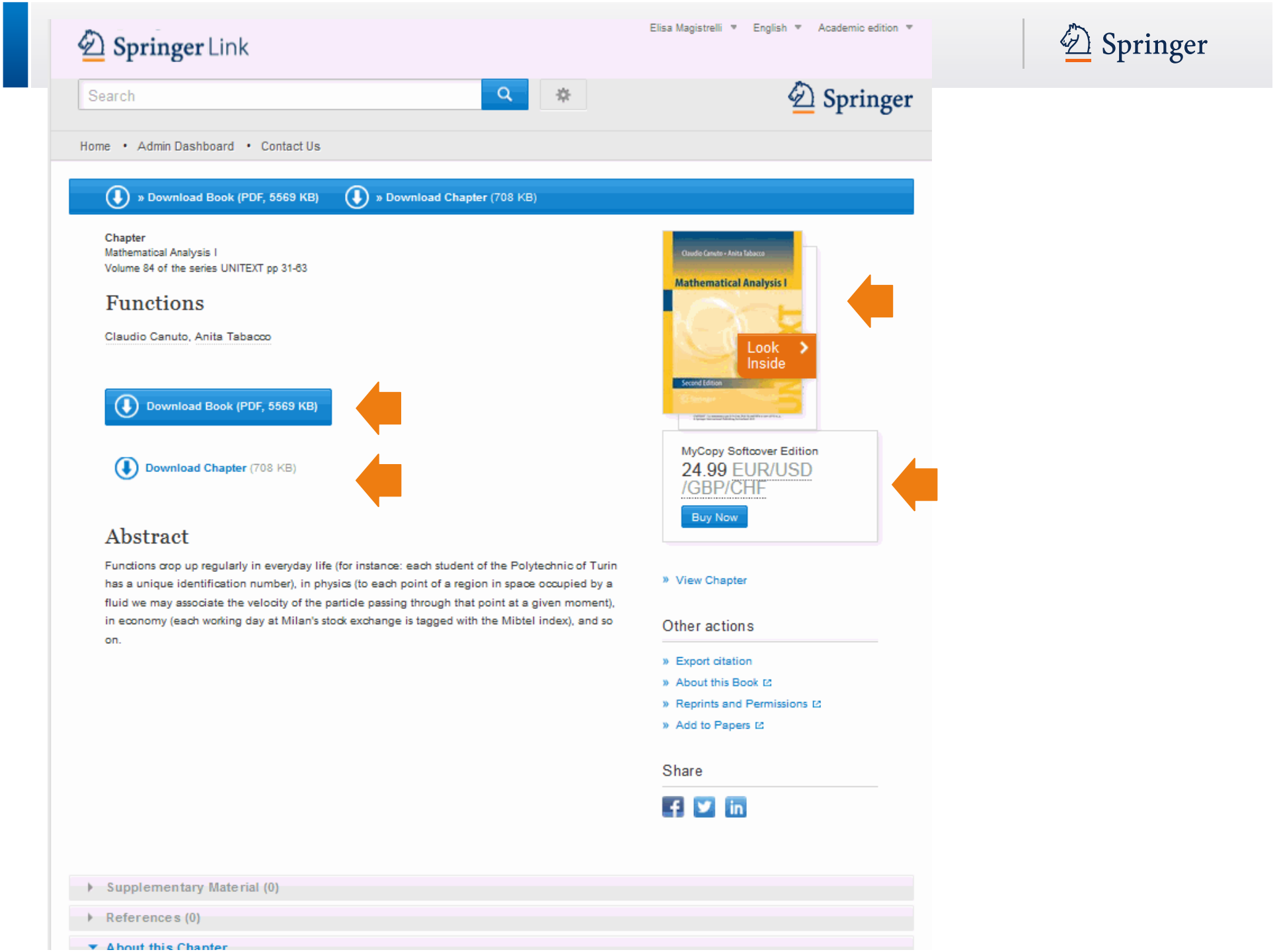
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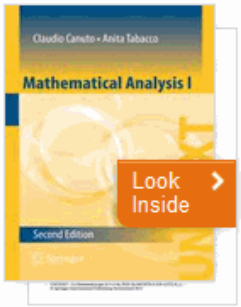

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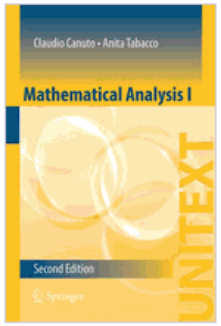

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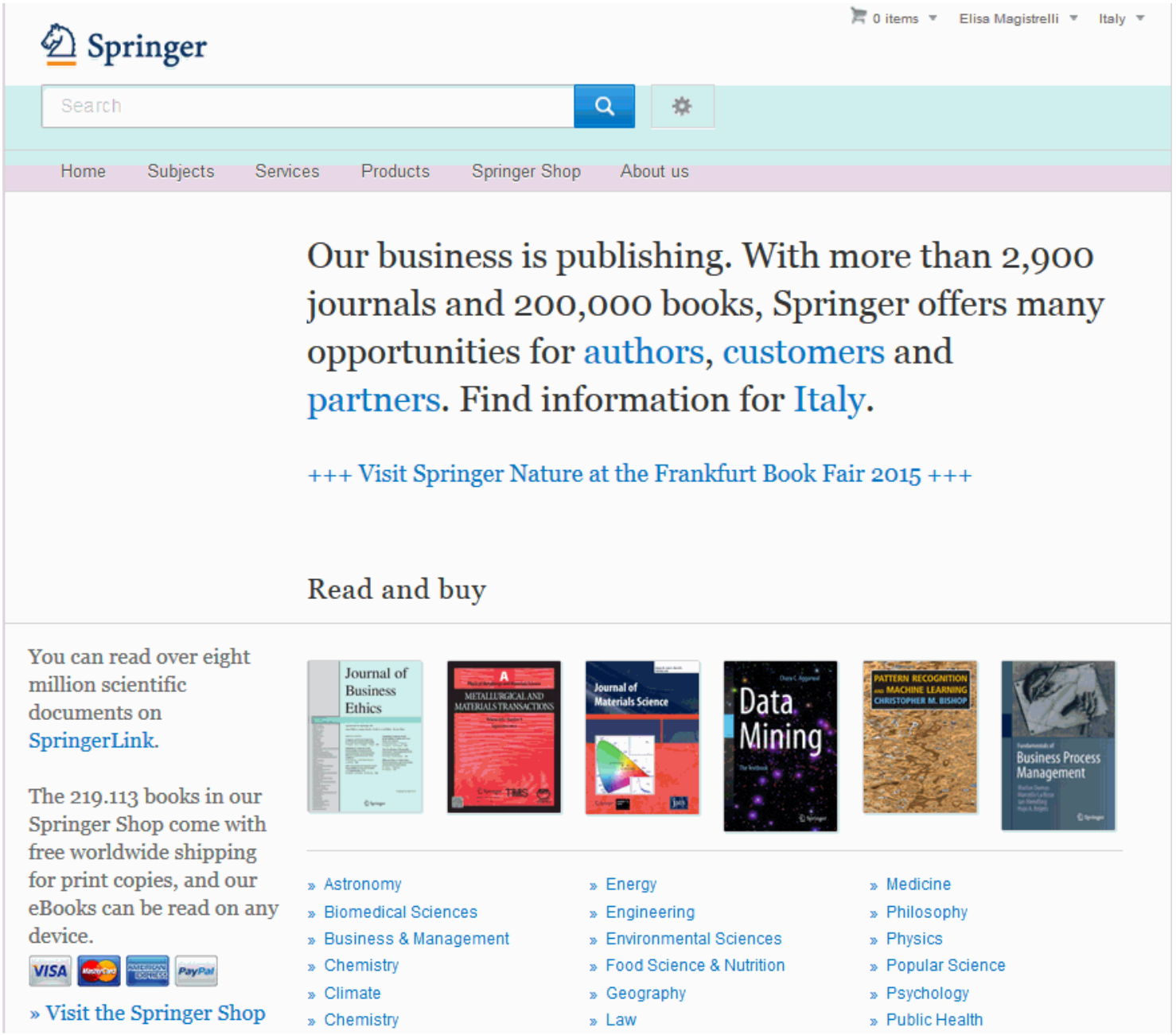

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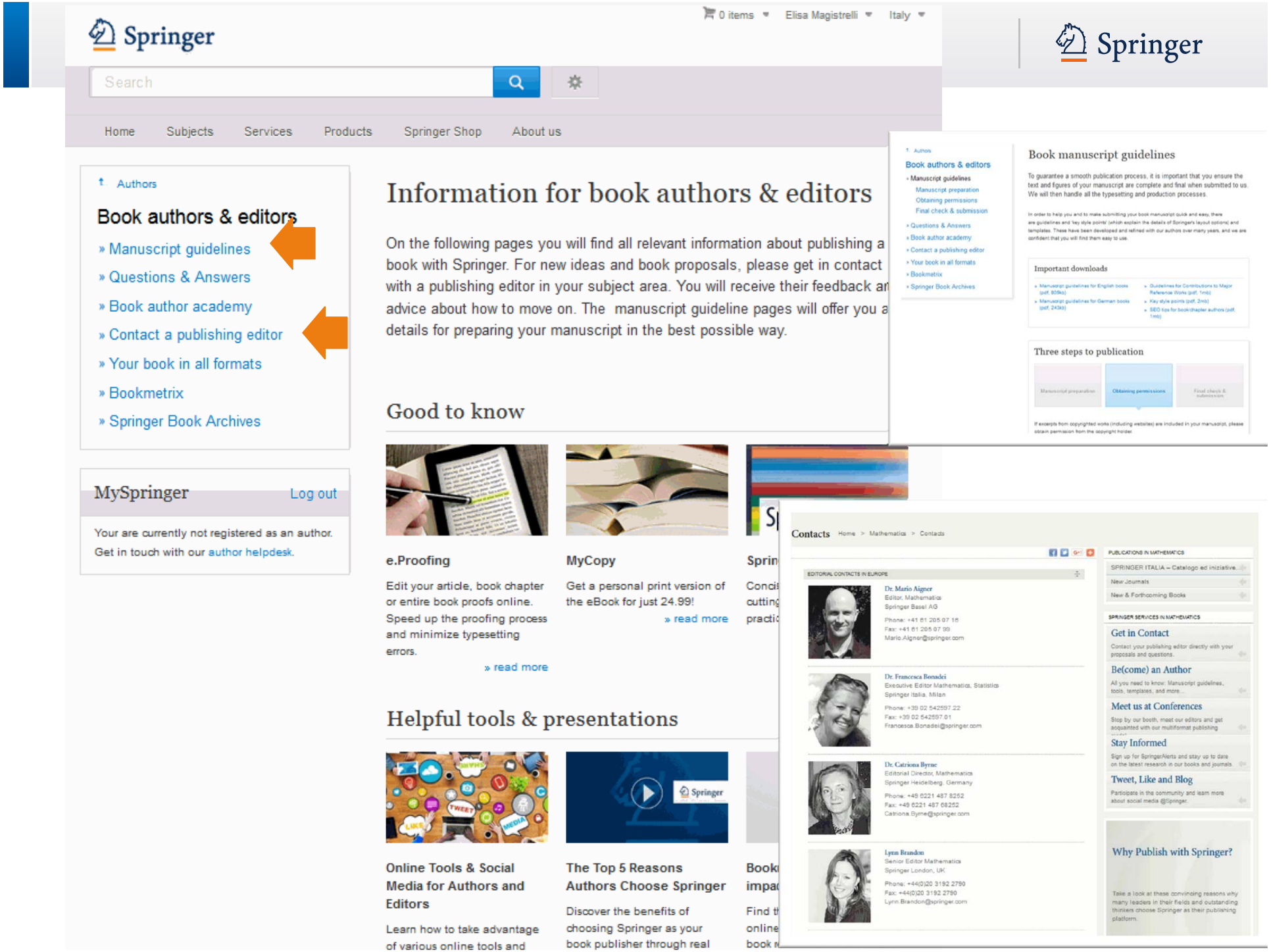
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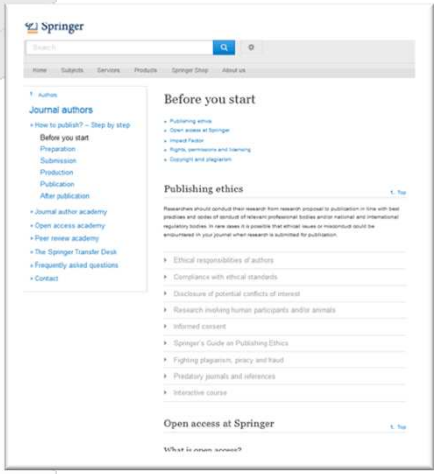

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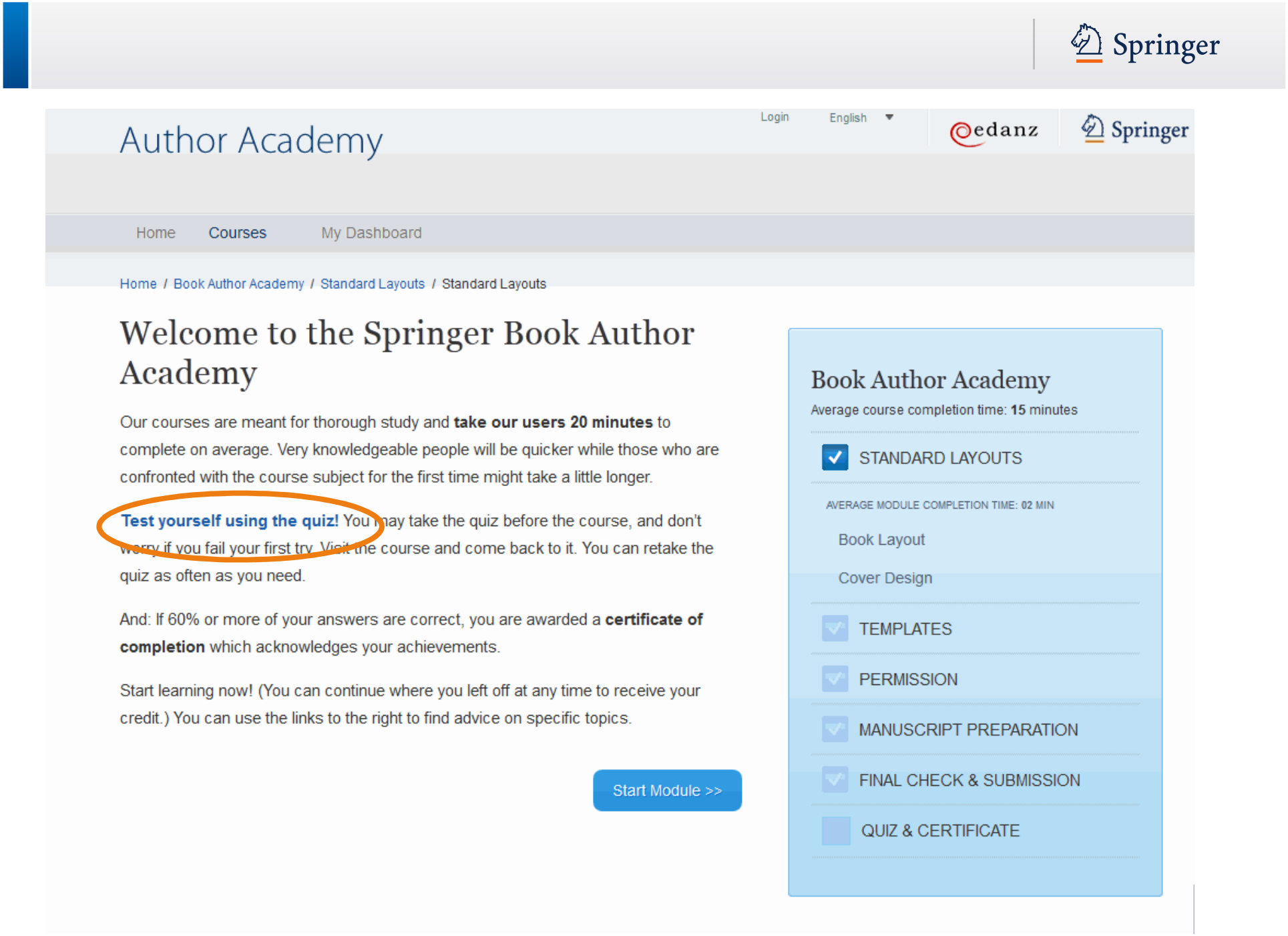
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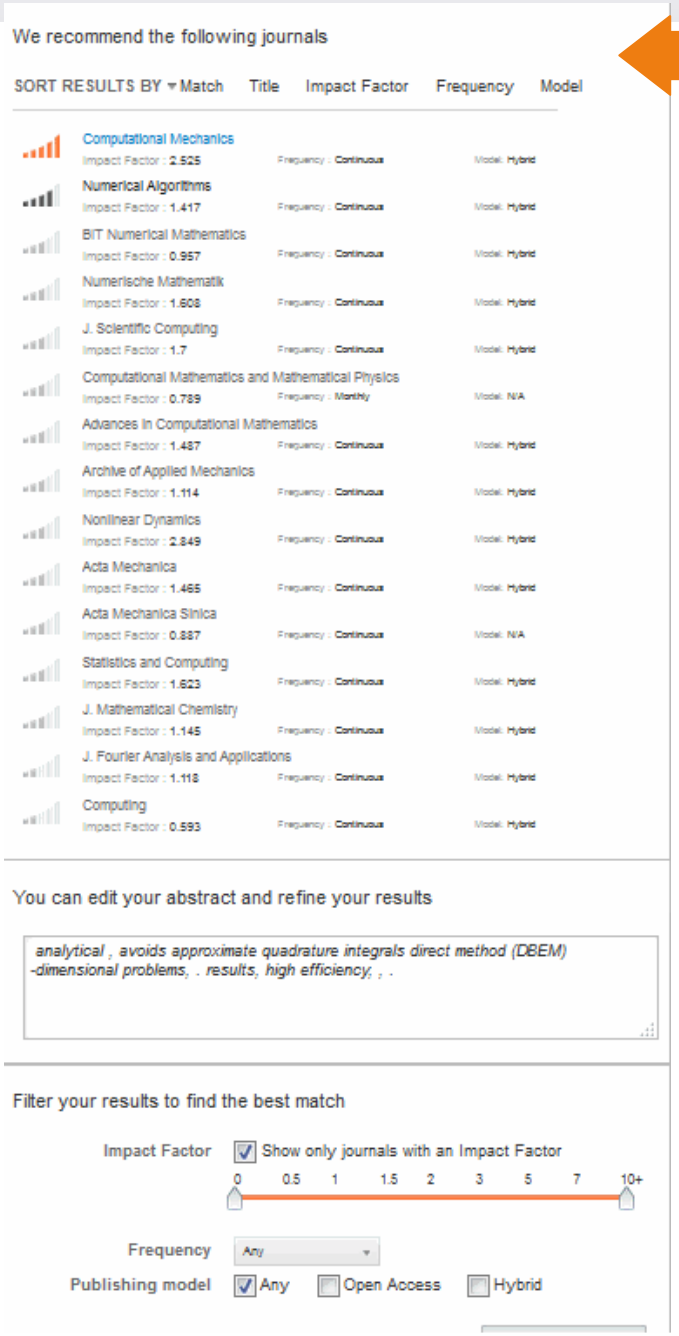
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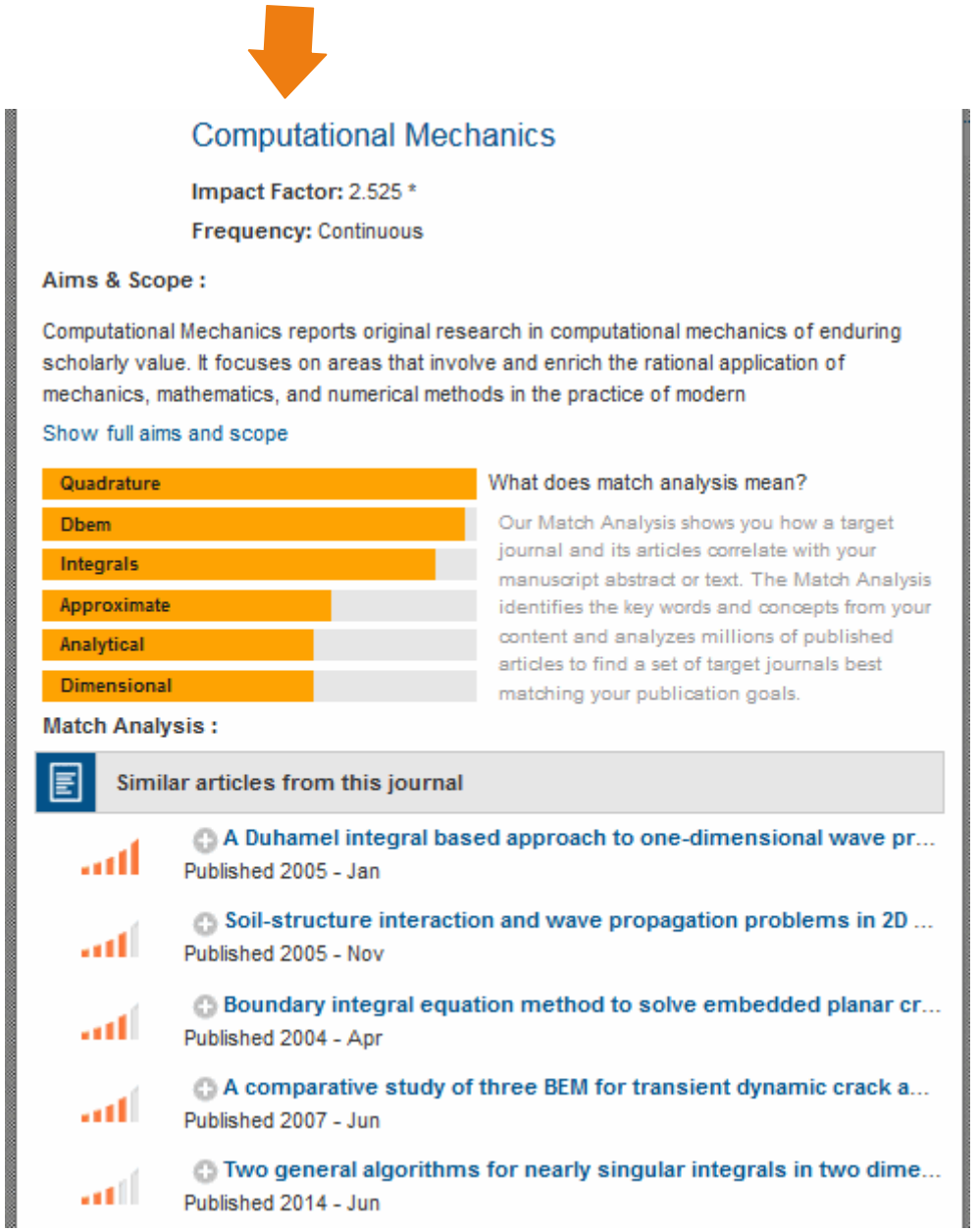
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The city of Thessaloniki is located in northern Greece, close to a large seismogenic area and it has experienced several destructive earthquakes during the present century. In this paper, we focus on the definition of the subsurface structure of the city from a site response analysis perspective. The paper presents, together with a summary of geology and seismicity, the results of a large-scale geophysical and geotechnical survey, in order to determine and validate geometry and dynamic properties of the main soil formations. The synthesis and combination of recent results regarding the dynamic properties with those obtained from the elaboration of a large database of classical geotechnical tests led to the design of a detail geotechnical map and of various 1-D profiles, 2-D cross sections and 3-D thematic maps for the main soil formations. These soil profiles and maps are oriented to site effect studies and provide a comprehensive picture easily adapted to geographic information systems (GIS) for planning and design purposes. Moreover, the results of this study were correlated with macroseismic observations reported in many earlier published microzonation studies of Thessaloniki. These comparative observations revealed the complexity of surface geology of the urban area, a fact which is expected to provoke additional amplification with respect to 1-D resonance.

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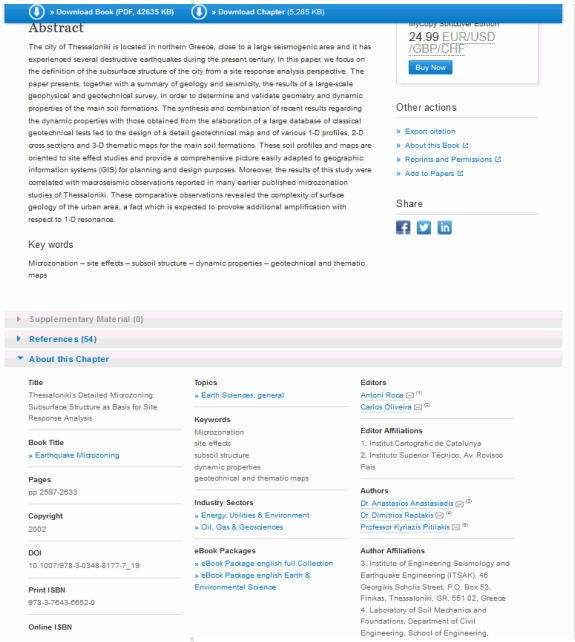

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Abstract

The city of Thessaloniki is located in northern Greece, close to a large seismogenic area and it has experienced several destructive earthquakes during the present century. In this paper, we focus on the definition of the subsurface structure of the city from a site response analysis perspective. The paper presents, together with a summary of geology and seismicity, the results of a large-scale geophysical and geotechnical survey, in order to determine and validate geometry and dynamic properties of the main soil formations. The synthesis and combination of recent results regarding the dynamic properties with those obtained from the elaboration of a large database of classical geotechnical tests led to the design of a detail geotechnical map and of various 1-D profiles, 2-D cross sections and 3-D thematic maps for the main soil formations. These soil profiles and maps are oriented to site effect studies and provide a comprehensive picture easily adapted to geographic information systems (GIS) for planning and design purposes. Moreover, the results of this study were correlated with macroseismic observations reported in many earlier published microzonation studies of Thessaloniki. These comparative observations revealed the complexity of surface geology of the urban area, a fact which is expected to provoke additional amplification with respect to 1-D resonance.

Key words

Microzonation – site effects – subsoil structure – dynamic properties – geotechnical and thematic maps

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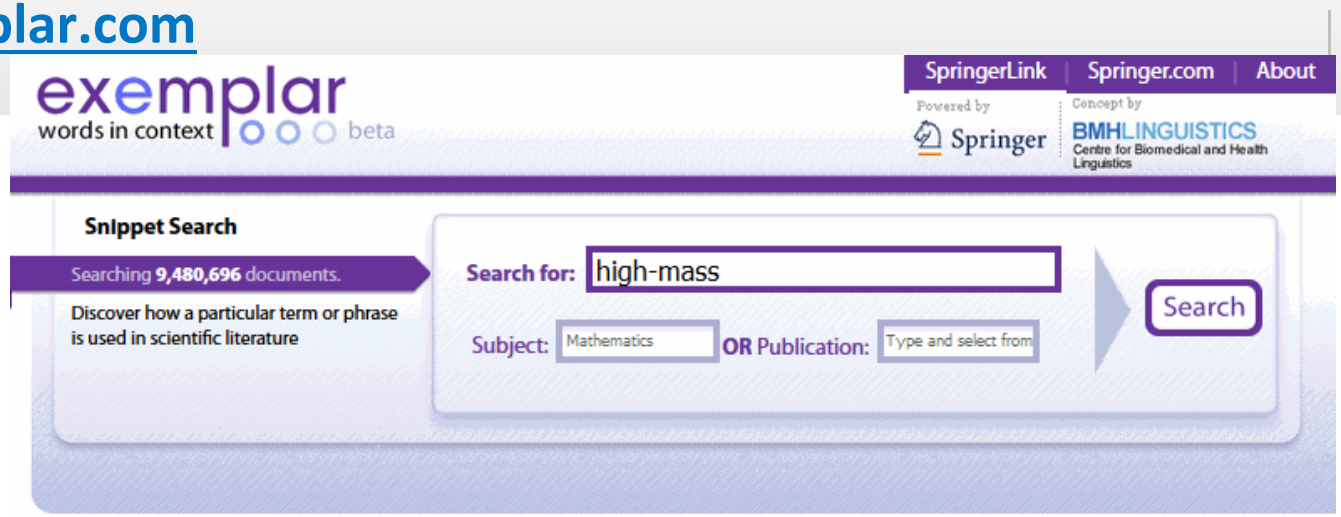

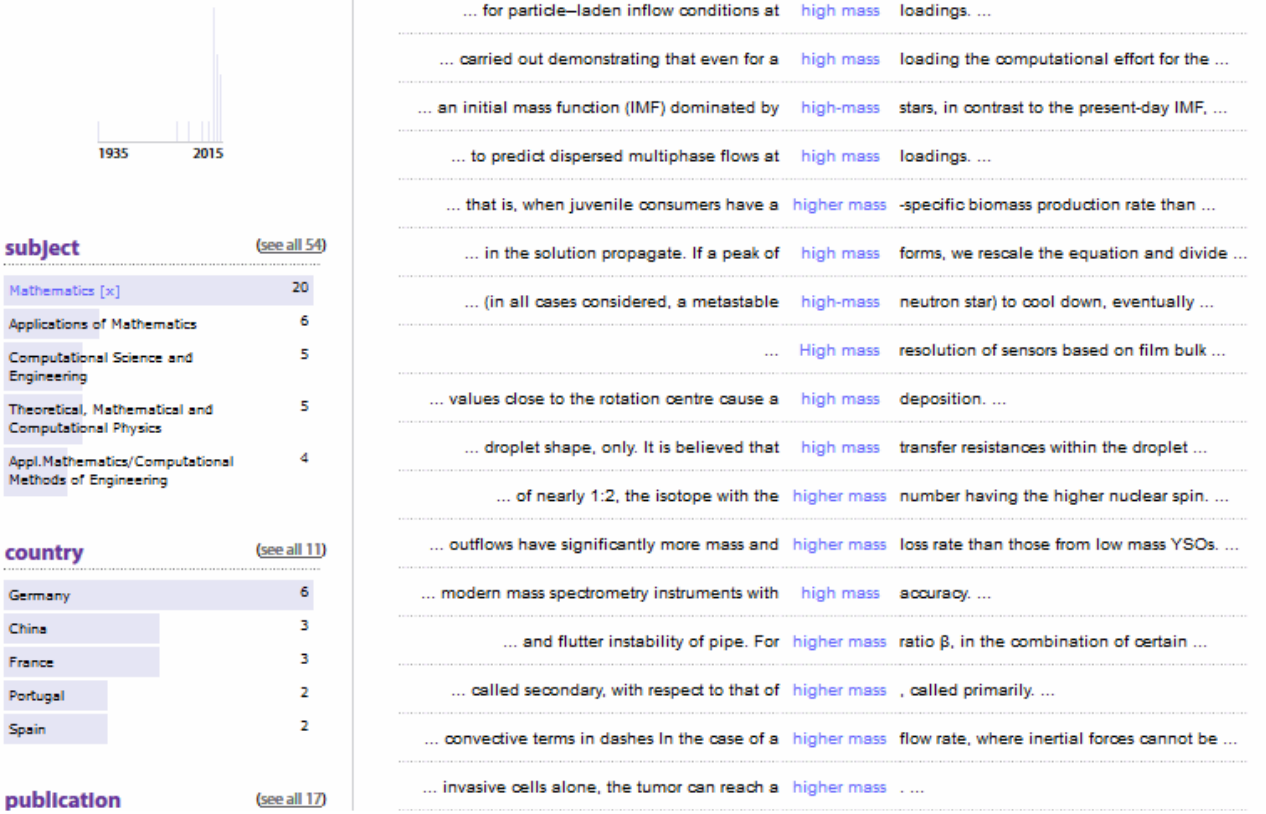

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... Fragmentation and Star Formation Because	high-mass	stars form deeply embedded in dense gas and ...
... 2002c). However, the stage prior to the	high mass	pr
... with the molecular and atomic gas.	High-mass	t
... model of structure formation implies that	high-mass	ga
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... its mass. The final stages of low-mass and	high-mass	stars are also different: low-mass stars ...
... The core masses in C1 and C2, and those in	high-mass	protostellar objects suggest an ...
... the observed power-law behavior at the	high-mass	end of the IMF, we need to add complexity ...
... of objects is larger the rarer they are.	High-mass	peaks exceeding the density threshold ...
... requires both high resolving power and	high mass	spectral accuracy to increa
... Galaxy opens a new door to understand how	high-mass	stars form in the context of
... for analysis of peptides and proteins with	high mass	accuracy and reduced bac
... MS n scans, and an orbitrap analyzer for	high mass	resolution and accurate ma
... ionization technique (IA-TOF). Although	high mass	resolution generally require
... regions using data-driven methods. In the	high-mass	region, additional decay mo
... an explanation as to why the b-ions in the	high mass	range are typically of lower
... to ToF, and benefit from intrinsic	higher mass	resolution and a reduction i
... ZZ^*$$ H → Z Z * decay at slightly	higher masses	. ...
... of approximately 700 cm, allowing both	high mass	resolution and high sensitivi
... has allowed to study IMF sampling at the	high mass	end and the concept of a cl
... heat sink. Such channels at relatively	high mass	flux have been studied [43,
... the phase space is generated in two parts:	high mass	"WZ " events with $$m_{{\e
... and (b) how cloud cores of both low and	high mass	that are super-critical subsequently ...

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CO, HI and Star Formation in the Large Magellanic Cloud

Yasuo Fukui

Abstract The Large Magellanic Cloud is an ideal laboratory for studying cloud evolution and star cluster formation. We will review CO and HI observations and compare observational signatures of star formation with the molecular and atomic gas. High-mass star formation shows good spatial correspondence with CO, and the three types, Types I, II and III of GMCs, with different star formation activity represent an evolutionary sequence from the initial inactive phase to the subsequent highly active phase in star formation. The dense clumps within GMCs show a similar evolutionary trend on smaller scales. It is likely that star formation takes place rapidly in 1–10 Myr roughly close to the free-fall timescale at 10–100 pc. GMCs are enveloped by relatively dense HI having density of 100 cm^{-3}. The envelope is continuously accreting onto the GMCs to increase the GMC mass with an average mass accretion rate of ~0.1 $M_{\odot}$ yr^{-1}. The HI filaments are formed by compression due to supergiant shells as seen in active cluster formation in 30 Dor, N159, N48, etc. Recent findings of the Galactic starburst regions suggest that the active high-mass star formation is often triggered by cloud-cloud collisions which cause efficient three dimensional compression of dense gas. The interface layer in such collisions creates higher turbulent speed and higher mass-accretion rate which facilitates rapid high-mass star formation in 10^5 yrs.

1 Introduction

Our scientific goal is to better understand evolution of galaxies. Because stars are formed in massive molecular clouds, formation and evolution of molecular clouds will provide key to elucidate formation of stars which build up galaxies. The Milky Way is not suited for such a study because of heavy contamination in the Galactic plane. On the other hand, the Large Magellanic Cloud (LMC) is an ideal object for such a study since it is located nearly face-on at a distance of 50 kpc. The LMC is also unique in studying high-mass star formation because young high-mass stars in populous clusters and in OB associations are extremely rich. In particular the

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\frac{{\alpha_{\gamma}}}
\tilde{\beta}
{\hbox{PE}}
"\epsilon" AND "2\pi"
"\bar \delta _q" OR "\frac{dw}{dz}"

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Self-intersections of random walks on lattices

Acta Mathematica Hungarica (2002) 96:187-220, August 01, 2002

$$P\left(E_n^{(d)}, \text{i.o.}\right) = 0 \quad \text{or} \quad 1$$

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P\left( {E_n^{\left( d \right)} ,{\text{i}}{\text{.o}}{\text{.}}}
\right) = 0\quad {\text{or}}\quad {\text{1}}
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$$\sqrt{s} = 800$$

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A measurement of the forward-backward asymmetry of e+e−→bb by applying a jet charge algorithm to lifetime tagged events

Zeitschrift für Physik C Particles and Fields (1995) 67:365-378, Septe

EXACT

$$A^b_{FB} = 0.062 \pm 0.034 \pm 0.002 - 0.082\Delta(\Gamma_{b\bar{b}}/\Gamma_{had})$$
$$at\sqrt{s} = 89.52GeV,$$
$$A^b_{FB} = 0.0963 \pm 0.0067 \pm 0.0038 - 0.471\Delta(\Gamma_{b\bar{b}}/\Gamma_{had})$$
$$at\sqrt{s} = 91.25GeV,$$
$$A^b_{FB} = 0.172 \pm 0.028 \pm 0.007 - 0.055\Delta(\Gamma_{b\bar{b}}/\Gamma_{had})$$
$$at\sqrt{s} = 92.94GeV,$$

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1 result

High-p T αα andαp interactions

Zeitschrift für Physik C Particles and Fields (1986) 30:513-520, Decen

EXACT

$$\sqrt{s} = 88$$

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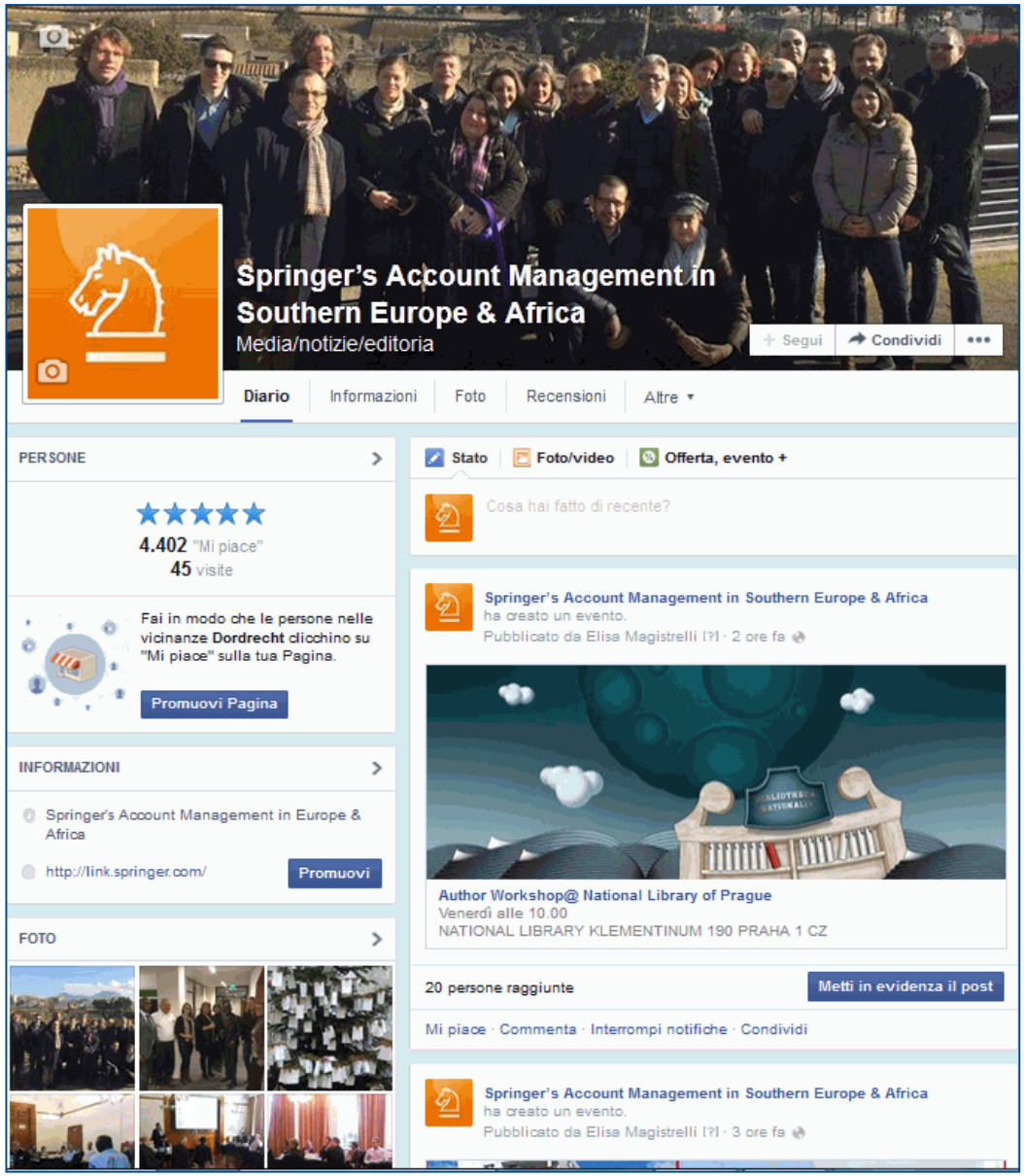

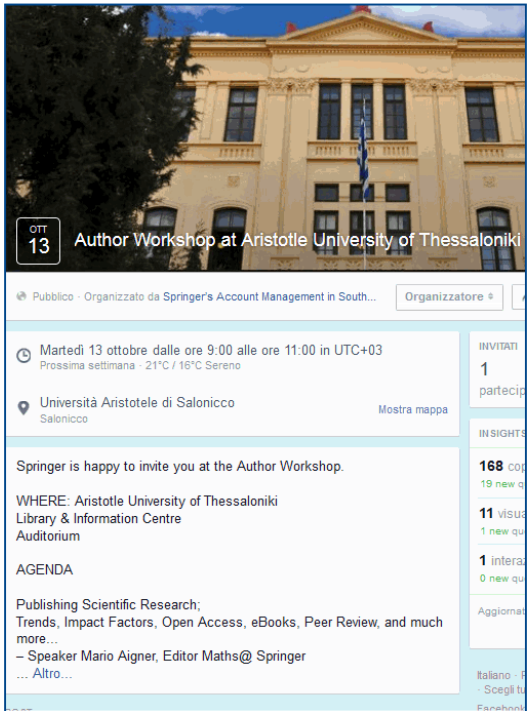

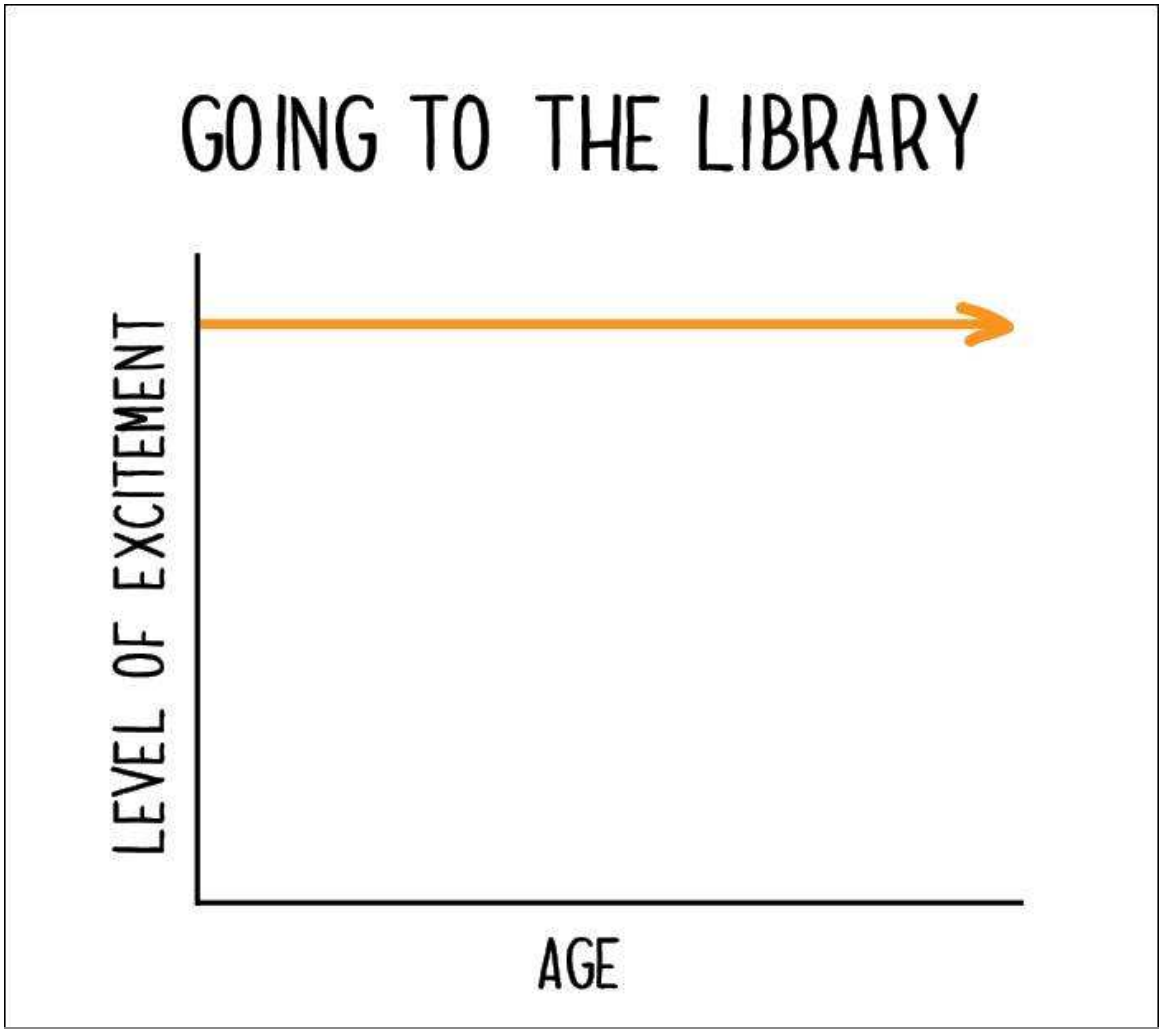

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